Measuring tangling in the solar photosphere

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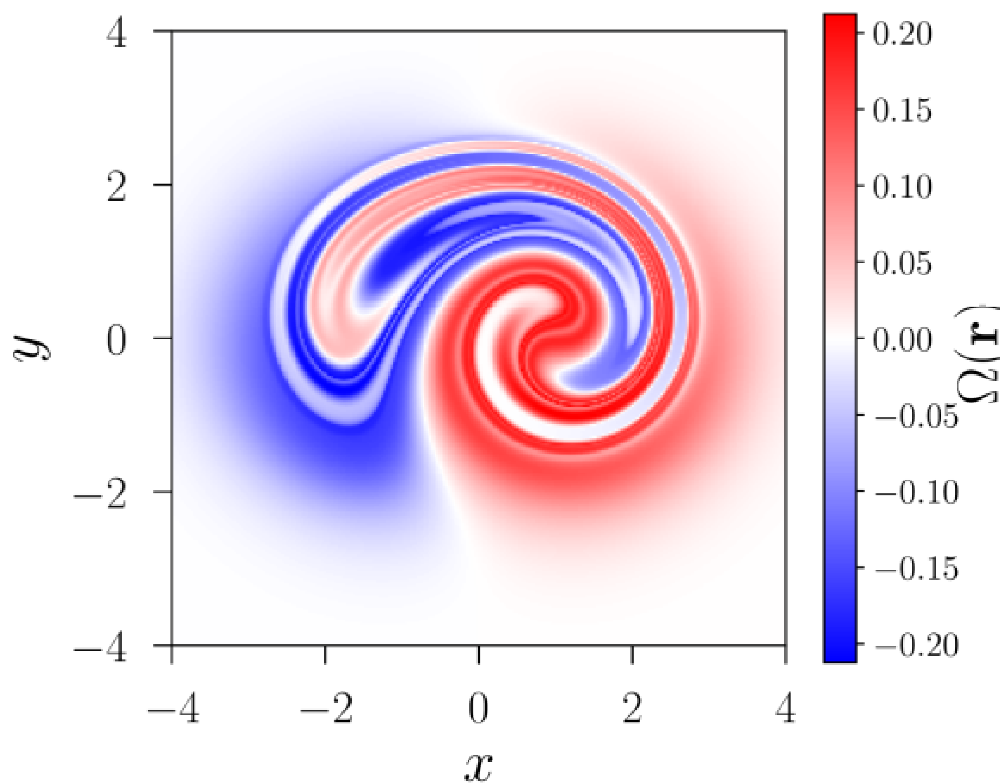
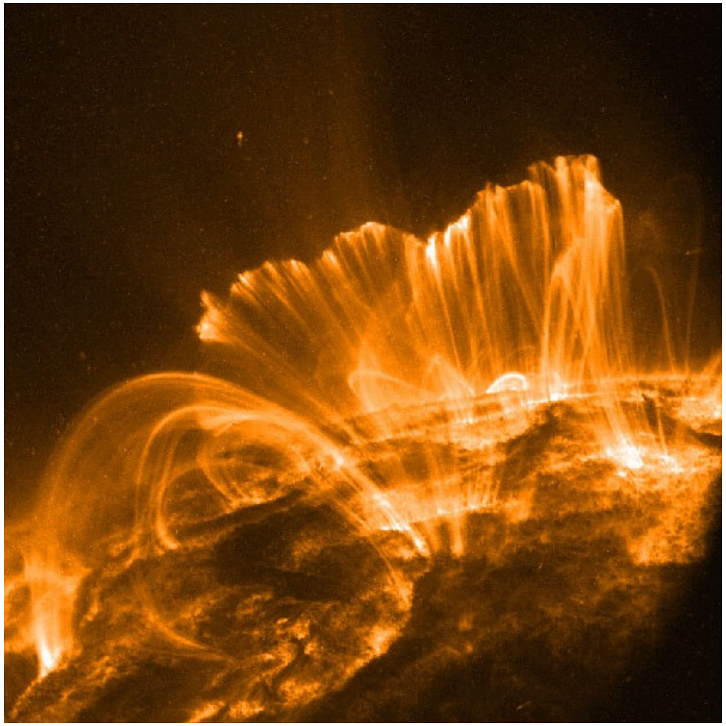
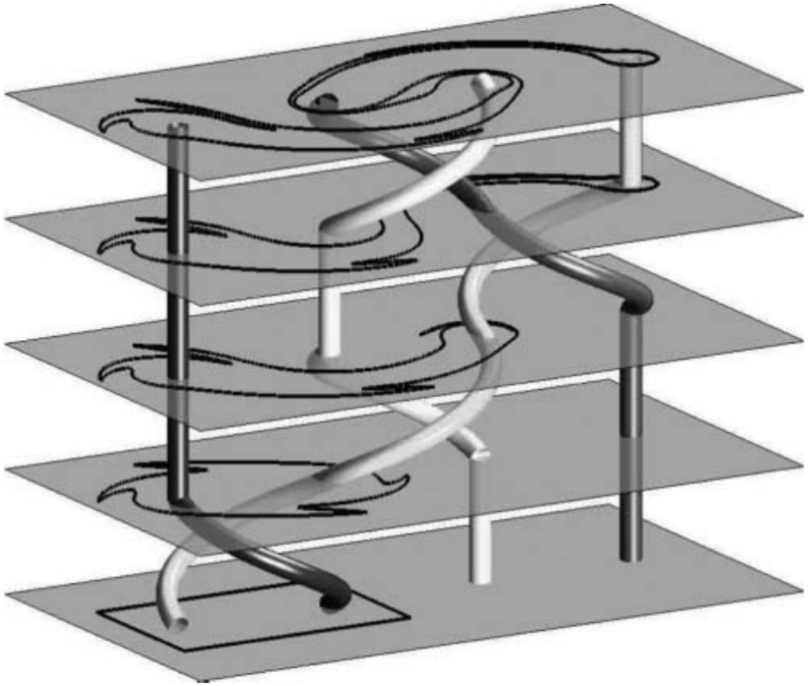
Anthony Yeates, Gunnar Hornig, Paul Bushby



Coronal Magnetic Fields

NASA

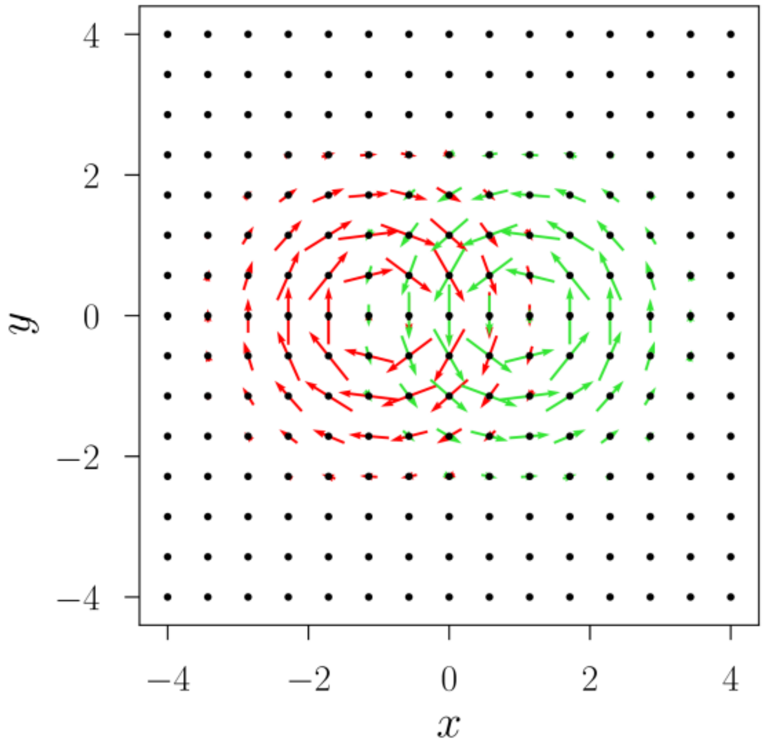


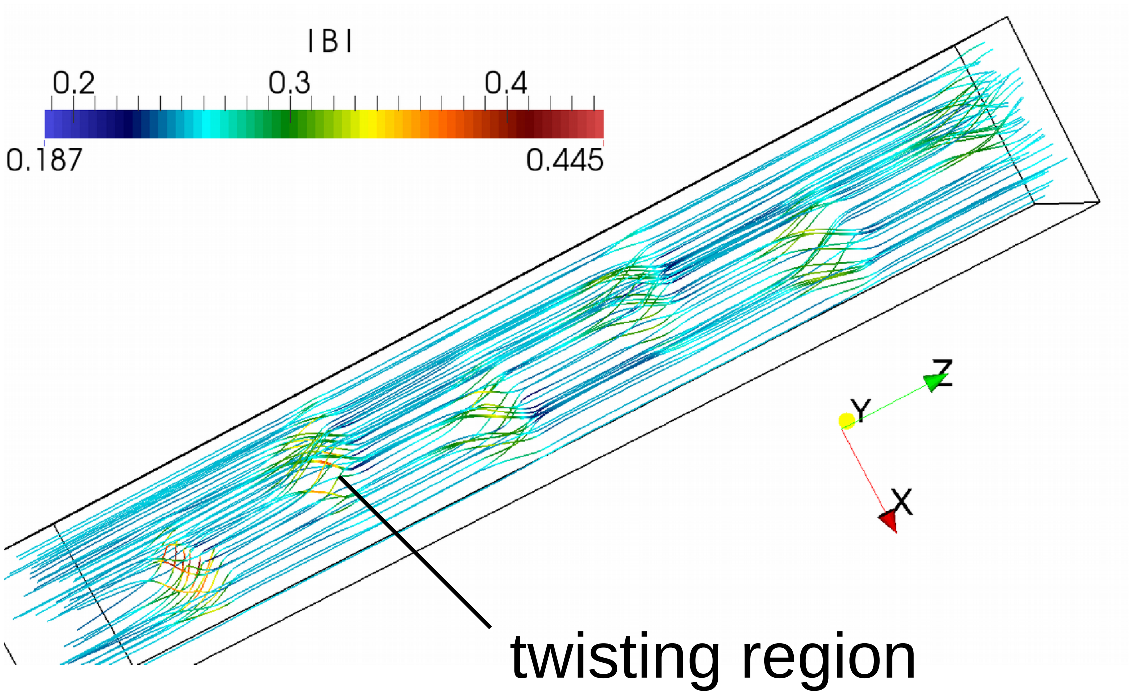


(Thiffeault et al. 2006)

- Field line tangling in solar magnetic fields.
- Study the tangling of solar magnetic field lines.

Blinking Vortex Benchmark

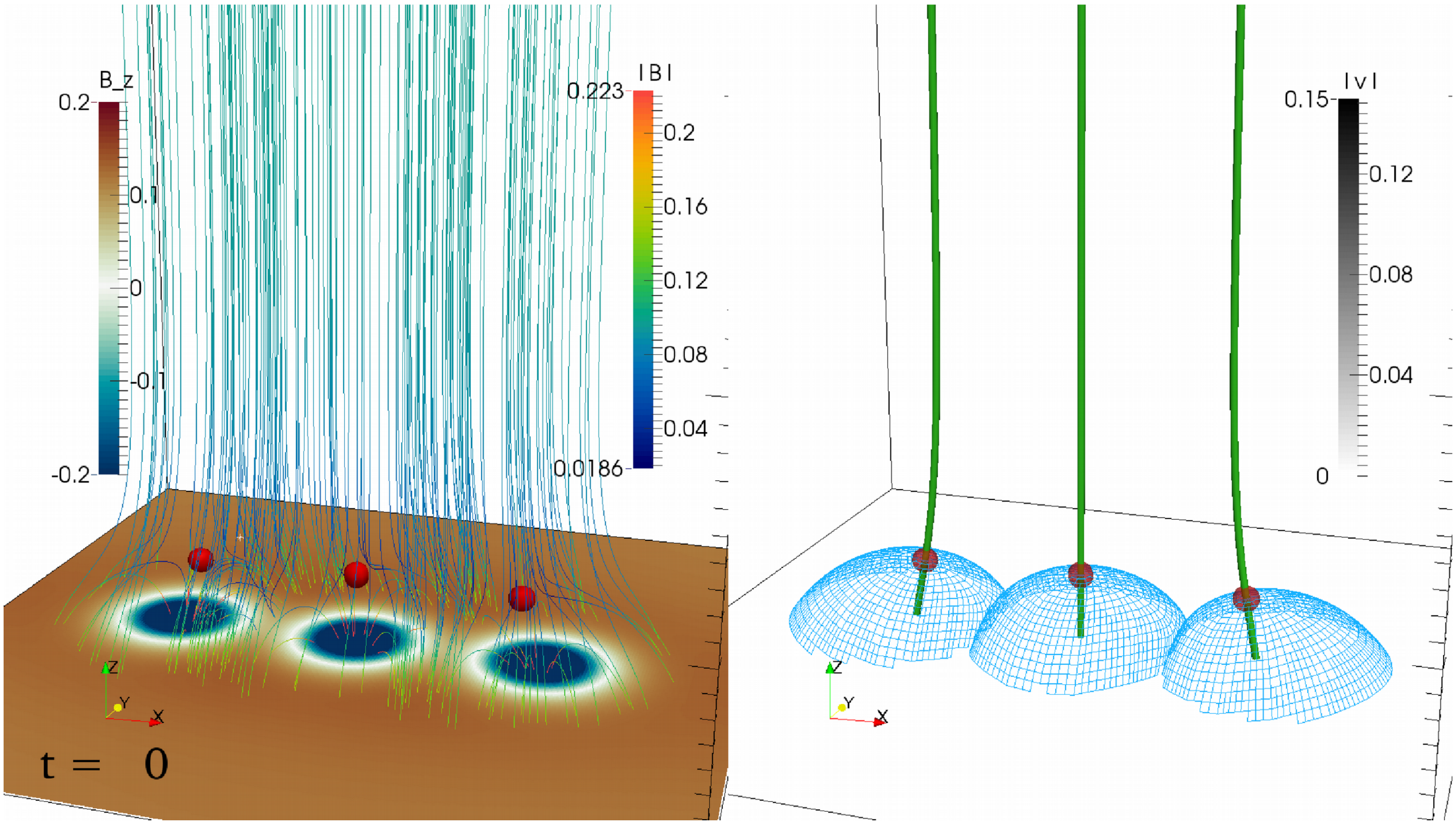




- Repeated applications of the blinking vortex motion.

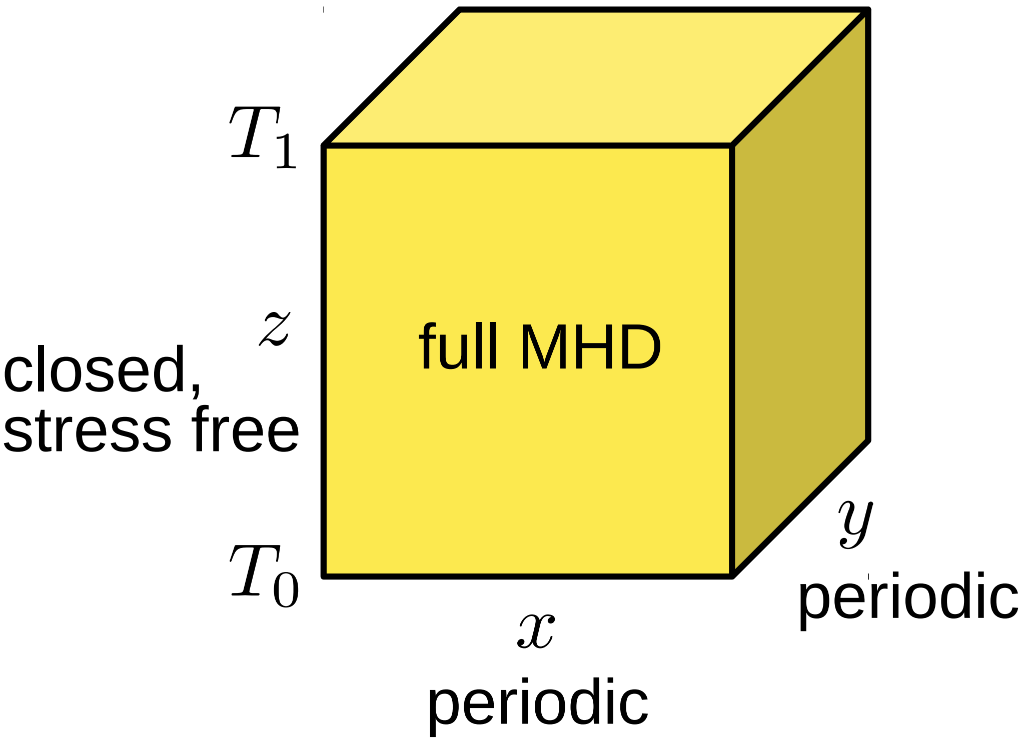
- World lines correspond to 3d braided magnetic field (pig tail, E3).

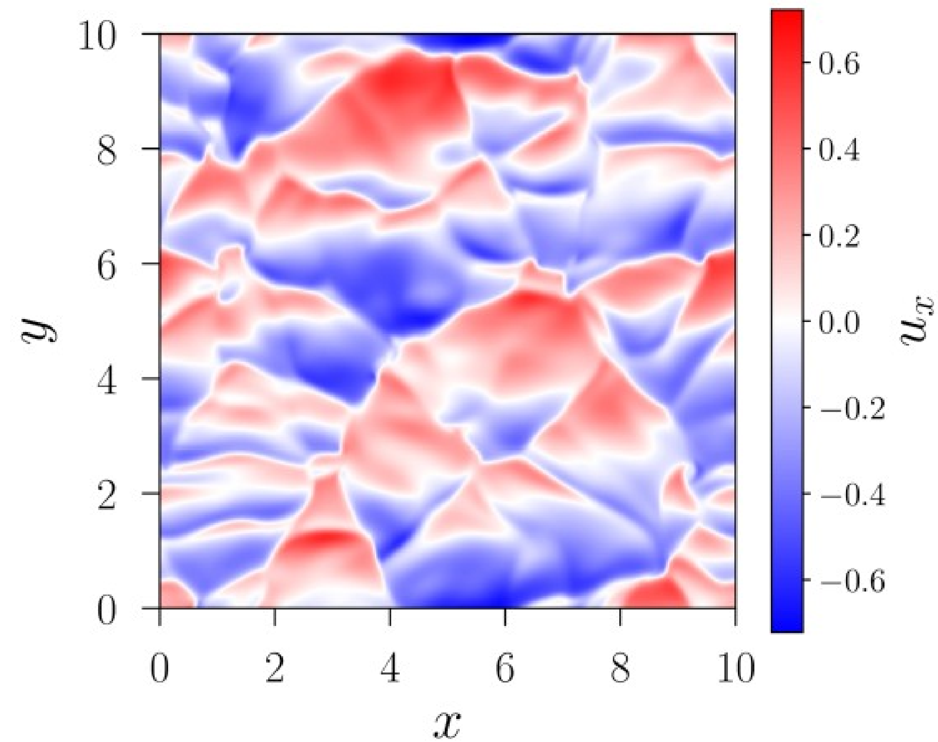
Driven Magnetic Fields in MHD



- Null pair creation/annihilation.
- Footpoint motion can alter the field line topology.

Magneto-Convection Simulations





(Bushby et al. 2012)

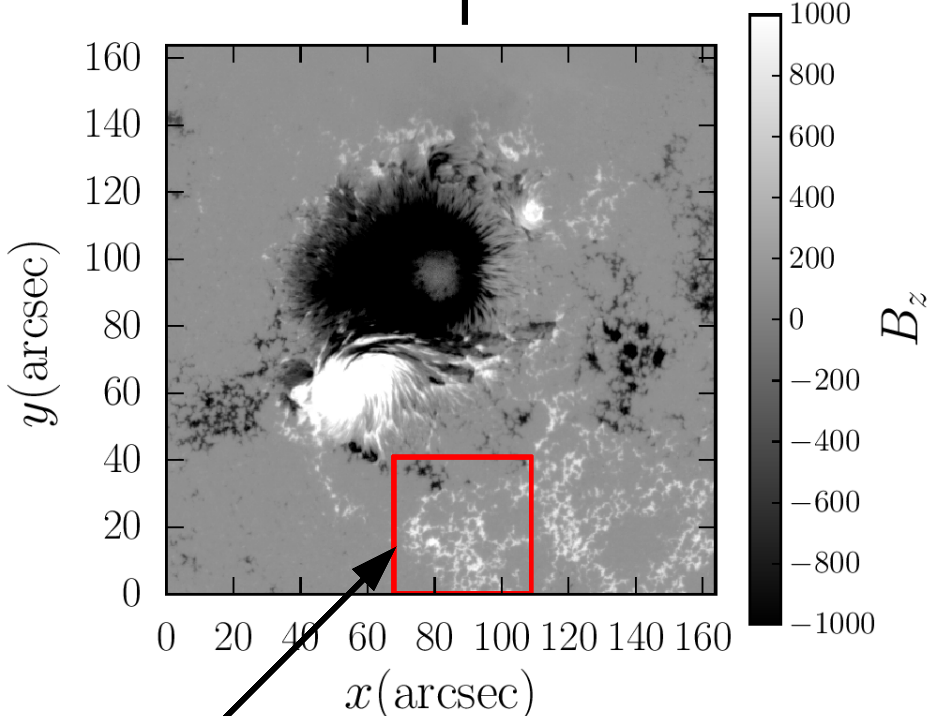
Helmholtz-Hodge Decomposition: $\mathbf{u} = \mathbf{u}_{\mathrm{i}} + \mathbf{u}_{\mathrm{c}} + \mathbf{u}_{\mathrm{h}}$

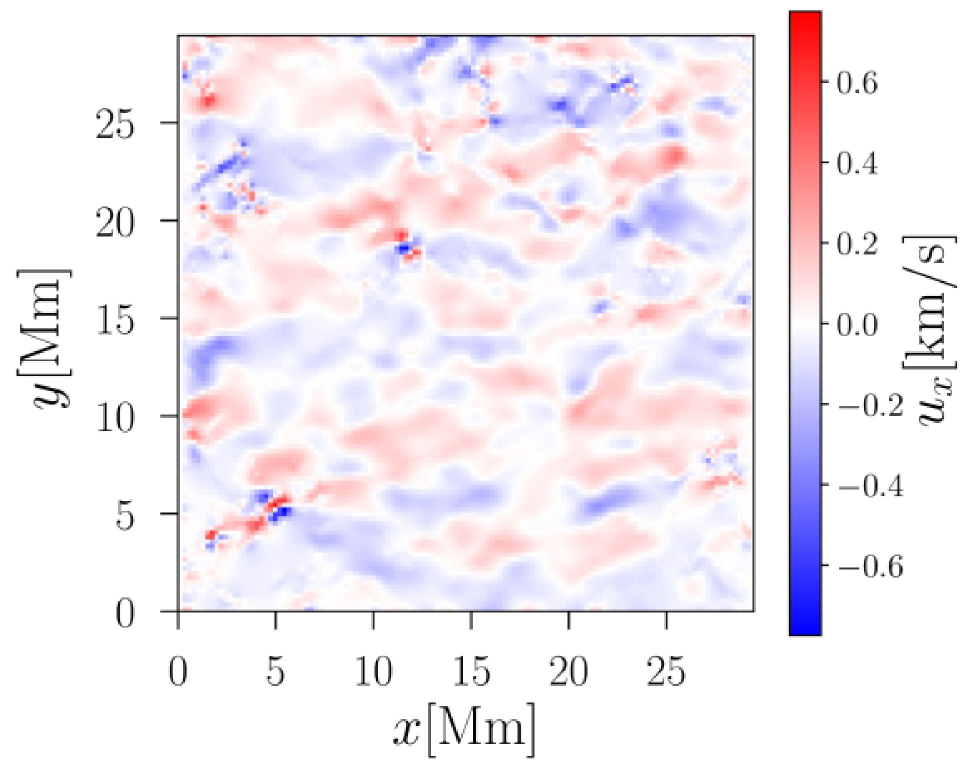
$$\mathbf{u}_{\mathrm{i}} = \nabla \times (\psi_z), \quad \mathbf{u}_{\mathrm{c}} = \nabla\phi, \quad \mathbf{u}_{\mathrm{h}} = \nabla\chi,$$

Active Region 10930

Helmholtz-Hodge Decomposition

FLCT

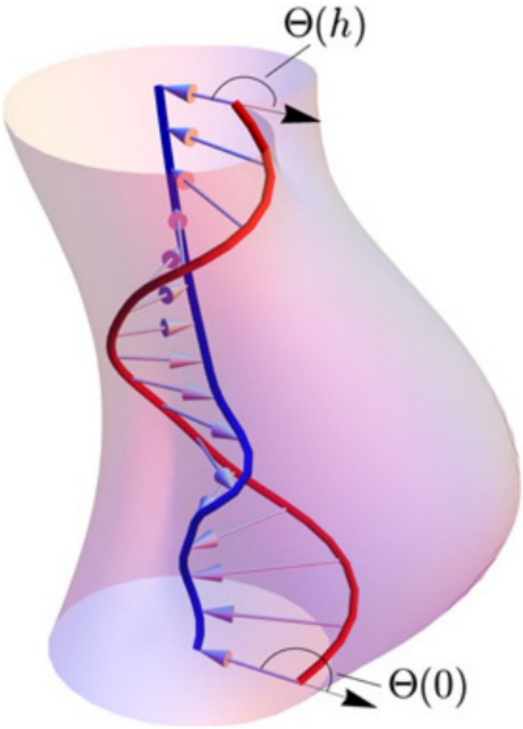




Consider this region.

12th of December 2006, 14:04 UT,
(Tsuneta et al. 2008, Fisher & Welsch 2008)

Winding Number



(Prior & Yeates 2014)

$$\frac{\mathrm{d}\mathbf{r}_1(t)}{\mathrm{d}t} = \mathbf{u}(\mathbf{r}_1(t), t) \qquad \frac{\mathrm{d}\mathbf{r}_2(t)}{\mathrm{d}t} = \mathbf{u}(\mathbf{r}_2(t), t)$$

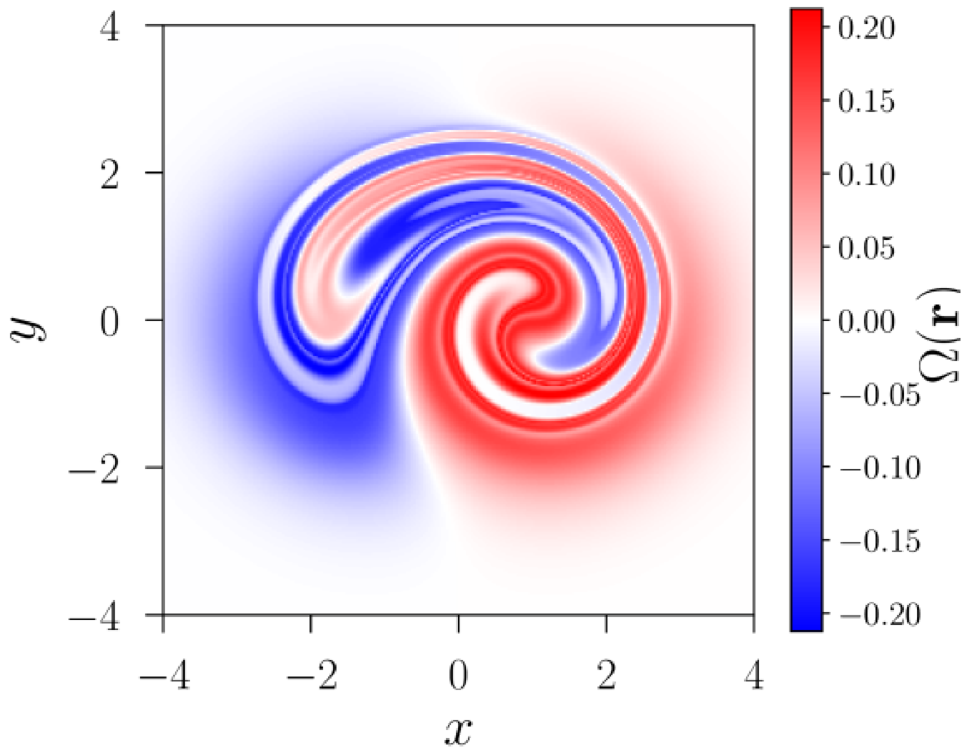
$$\Theta(\mathbf{r}_1, \mathbf{r}_2, t) = \arctan\left(\frac{y_2(t) - y_1(t)}{x_2(t) - x_1(t)}\right)$$

$$\Theta(\mathbf{r}_1, T) = \frac{1}{L_x L_y} \int_0^T \int_{(0,0)}^{(L_x, L_y)} \frac{\mathrm{d}\Theta(\mathbf{r}_1, \mathbf{r}_2, t)}{\mathrm{d}t} \, \mathrm{d}\mathbf{r}_2 \, \mathrm{d}t$$

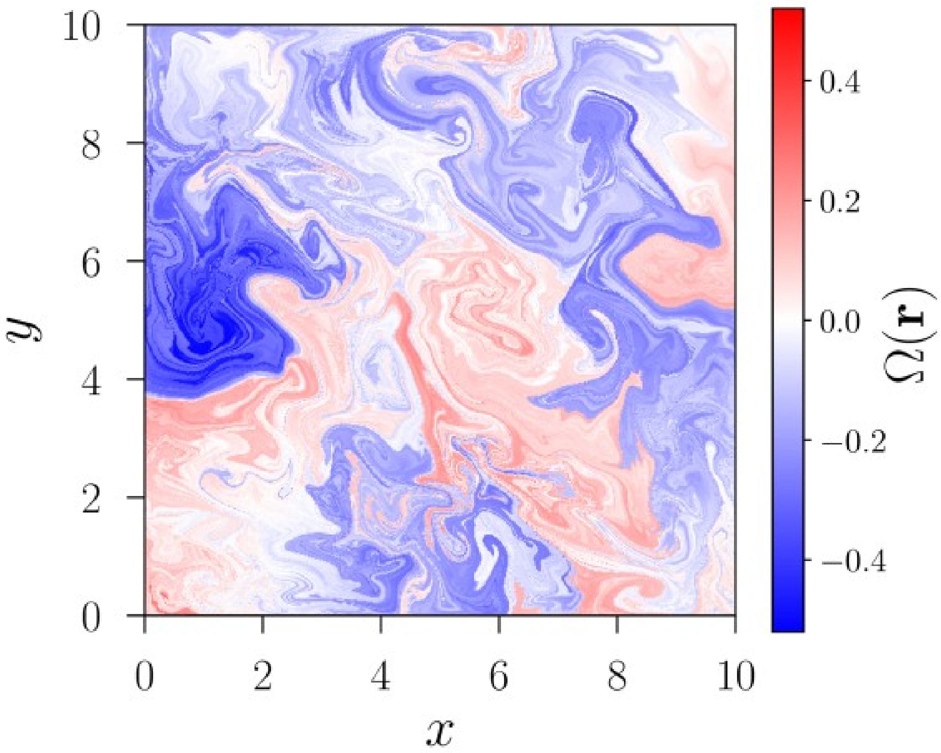
normalized averaged winding number:

$$\Omega(\mathbf{r}_1, T) = \frac{\Theta(\mathbf{r}_1, T)}{q(T)}$$

Winding Number

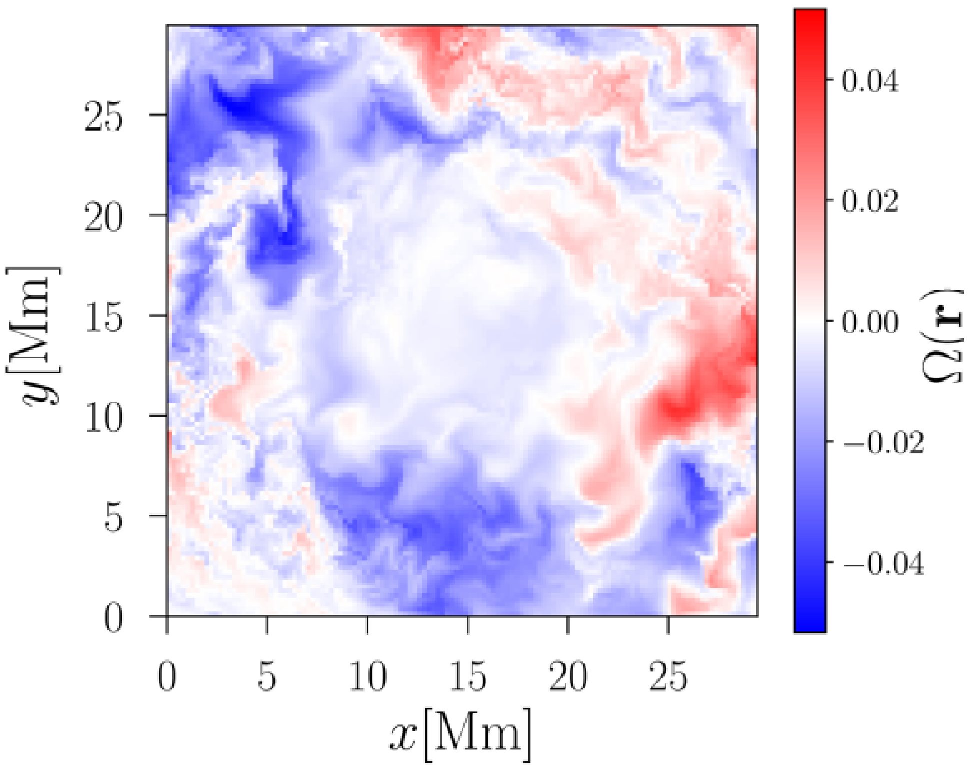


blinking vortex



simulation

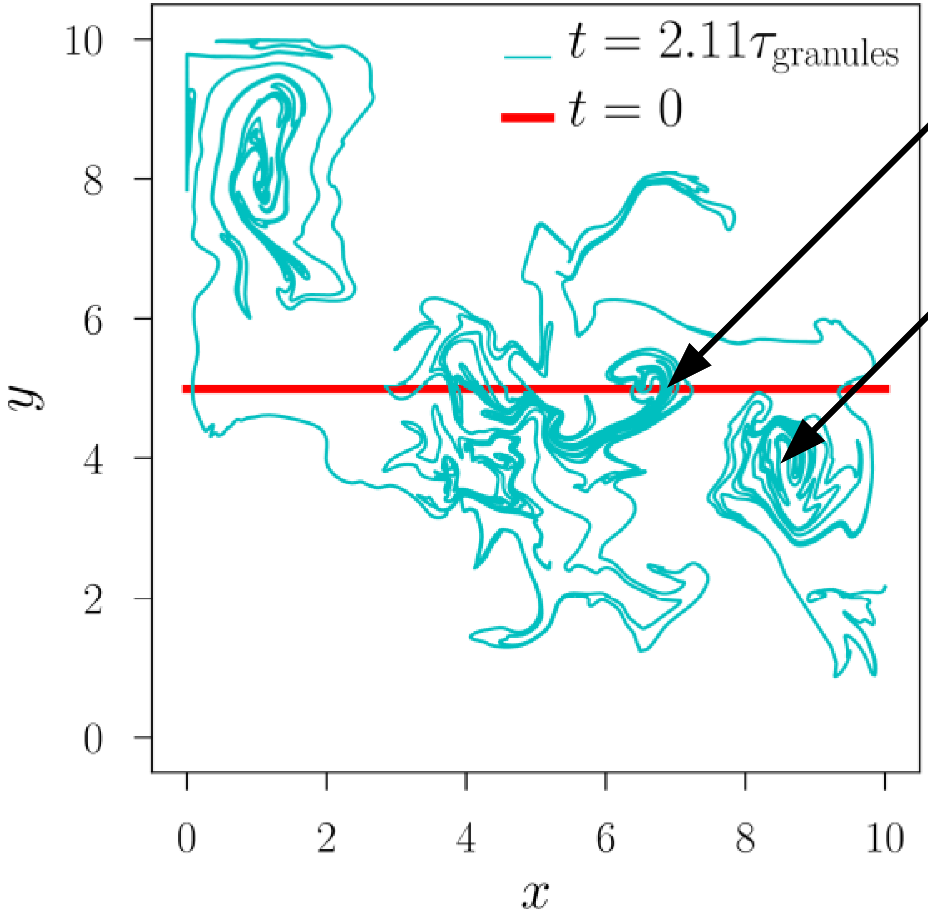
Winding Number



observations

- High winding for simulations.
- Low winding for observations.

Finite Time Topological Entropy



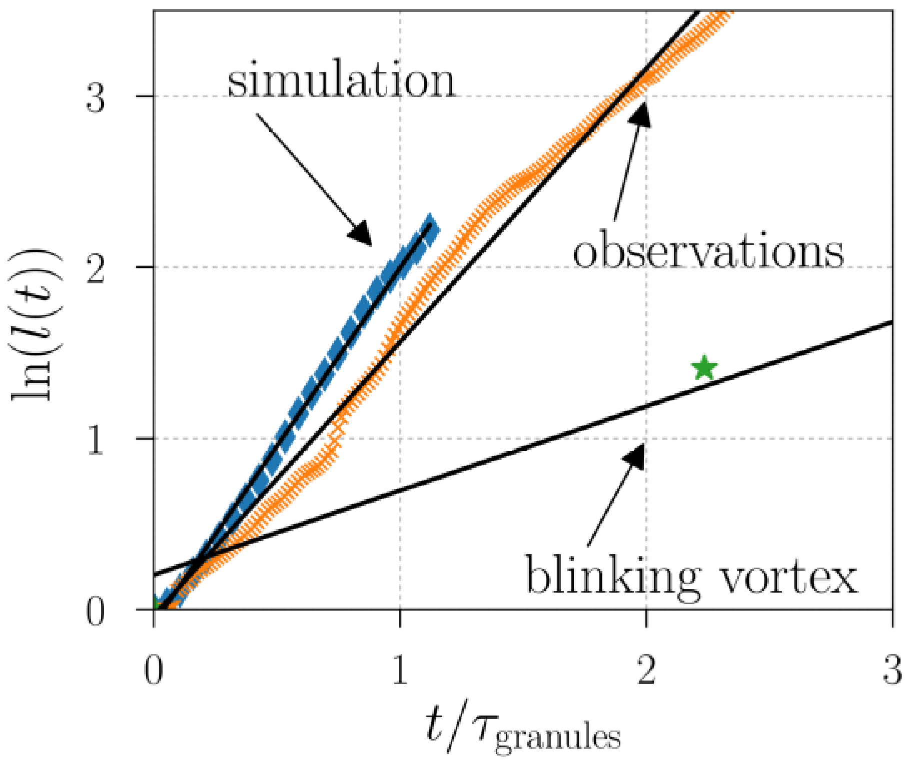
material line γ

advected material line $F(\gamma)$

FTTE:

$$h(F, \gamma, t) = \frac{1}{t} \ln \left(\frac{l(t)}{l_0} \right)$$

Finite Time Topological Entropy



- High tangling for simulations and observations.
- It takes 3.059h for the photosphere to get as tangled as during for one cycle of the blinking vortex motion.

Conclusions

- Driving changes magnetic field topology
- High degree of winding possible.
- High degree of entanglement
- Tangled magnetic field stores free energy to be released in reconnection events.

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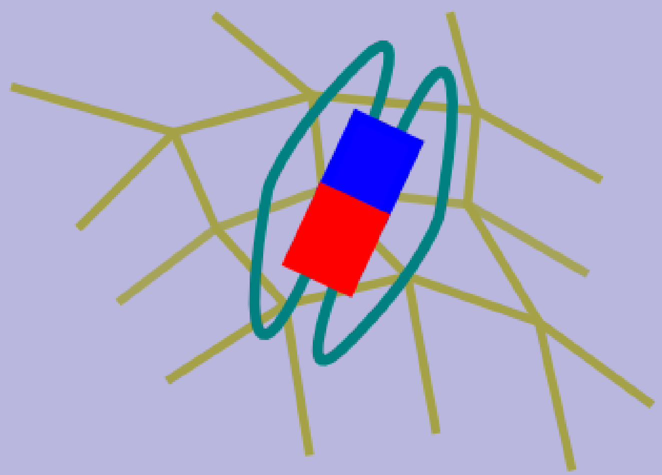
Numerical Methods in MHD

Scottish Numerical Methods Network 2018

7 September 2018

University of Dundee



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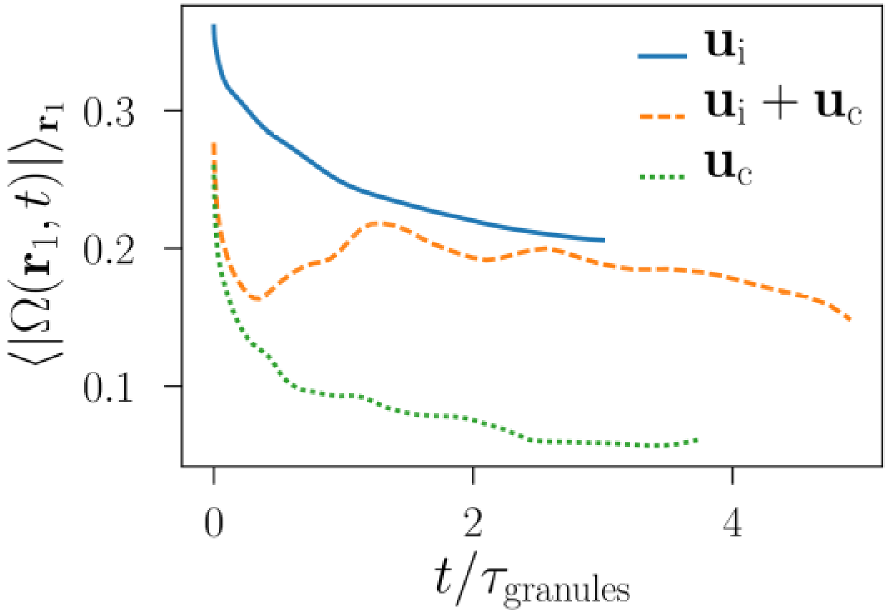
Winding Number

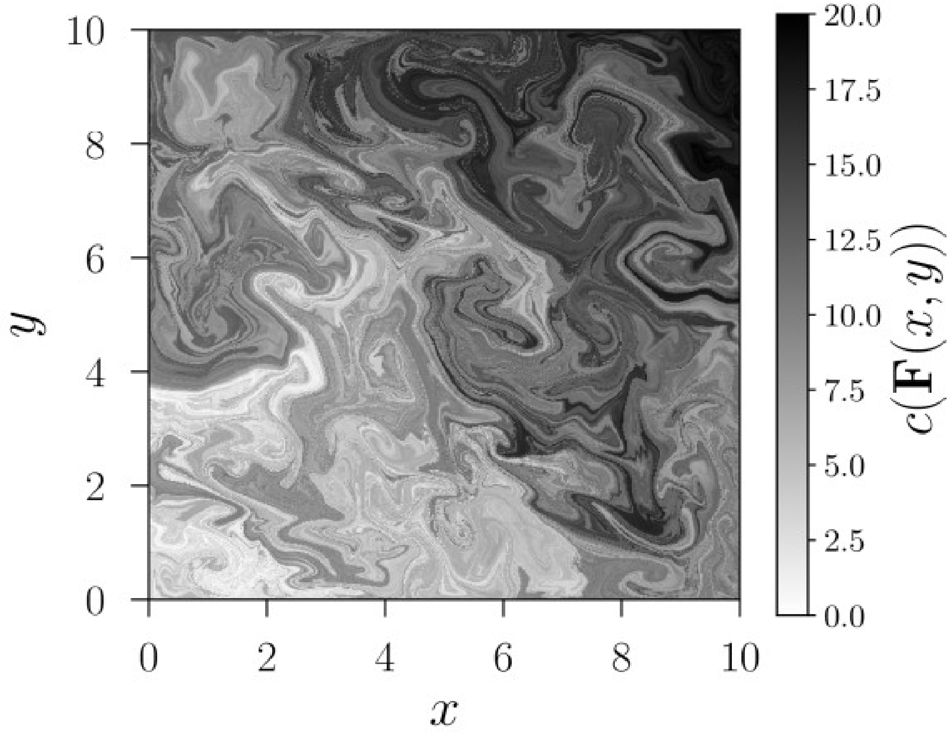
normalization: $$q(T) = \frac{1}{l_{\text{granules}} L_x L_y} \int_0^T |\mathbf{u}| \; \mathrm{d}x \; \mathrm{d}y \; \mathrm{d}t.$$

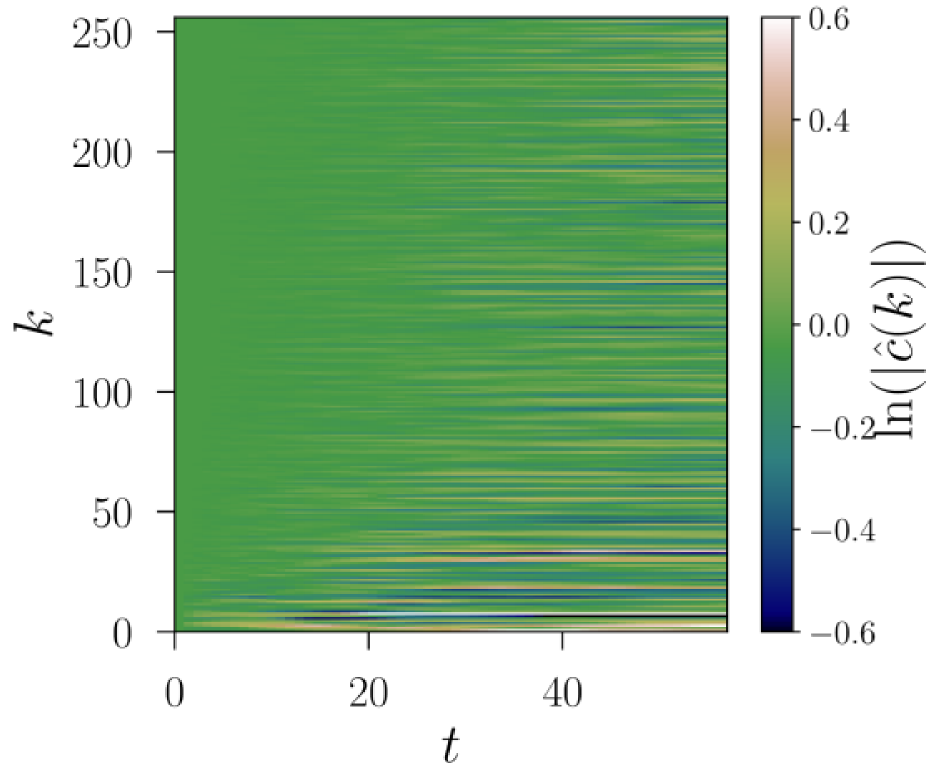


$$\mathbf{u} = \mathbf{u}_\mathrm{i} + \mathbf{u}_\mathrm{c} + \mathbf{u}_\mathrm{h}$$

Compressional part does not significantly contribute to the winding.

Passive Scalar





initial profile: $c(x, y) = x + y$

- High mixing of passive scalar.
- No clear scale due to turbulent motions.